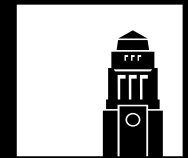


GAMMA-RAY ASTRONOMY AT THE HIGHEST ENERGIES WITH CTA

RICHARD WHITE

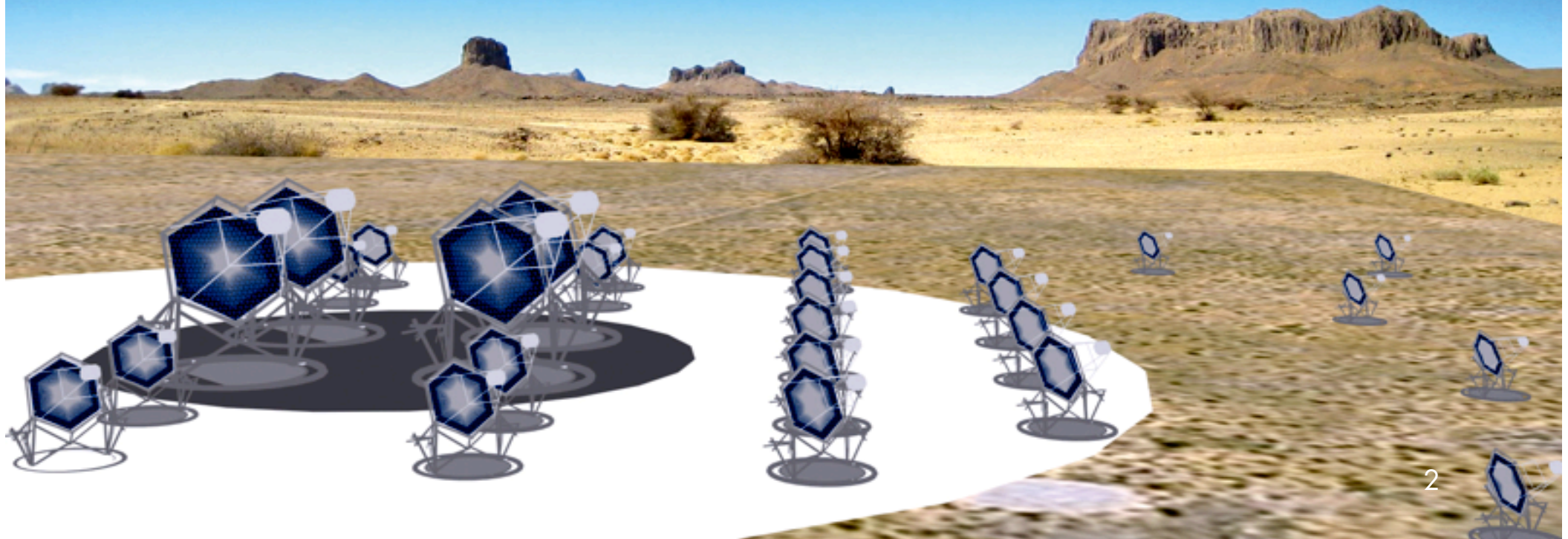
FOR THE UK CTA GROUPS

IOF, UCL, MARCH 2010

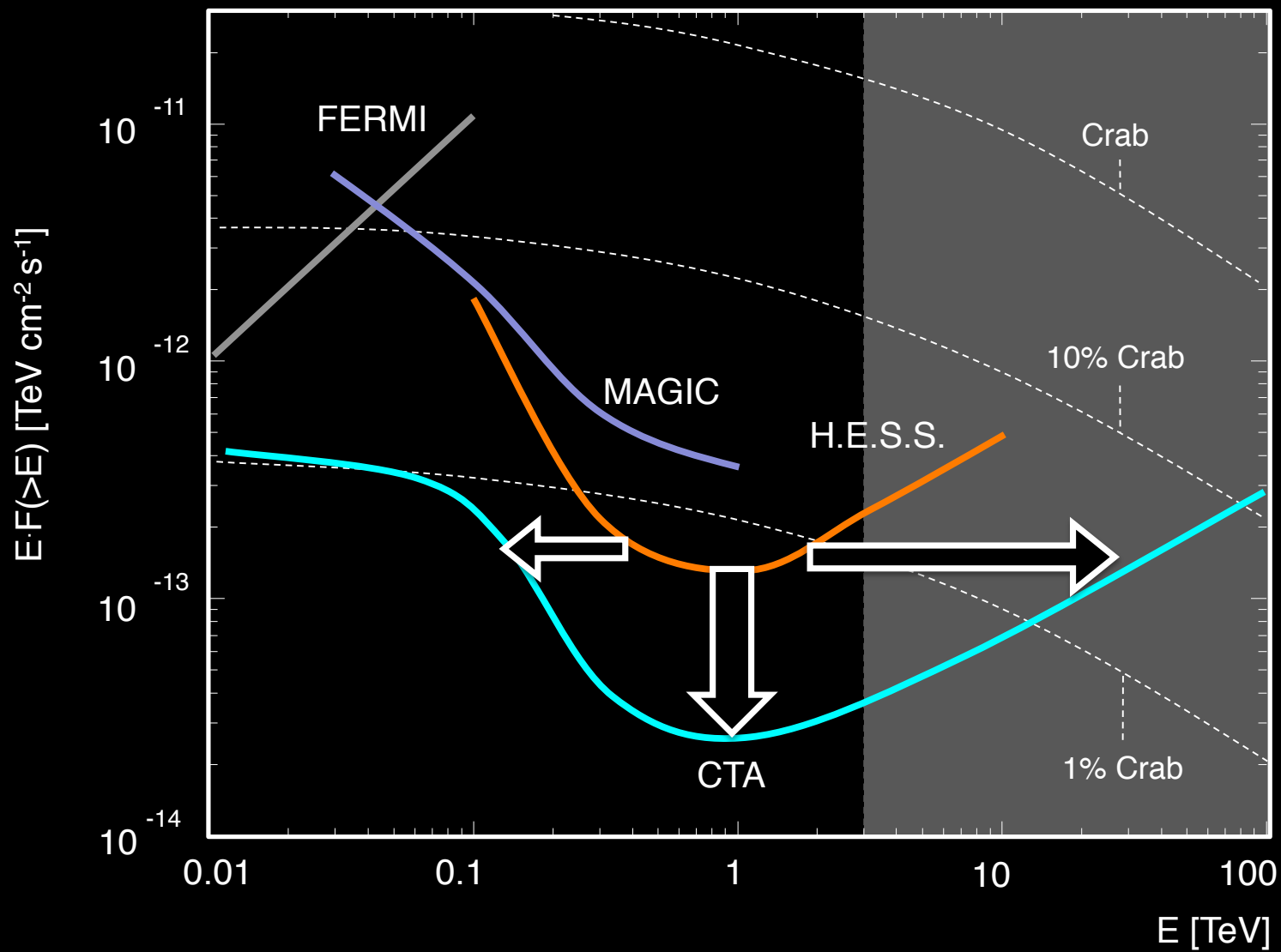


UNIVERSITY OF LEEDS

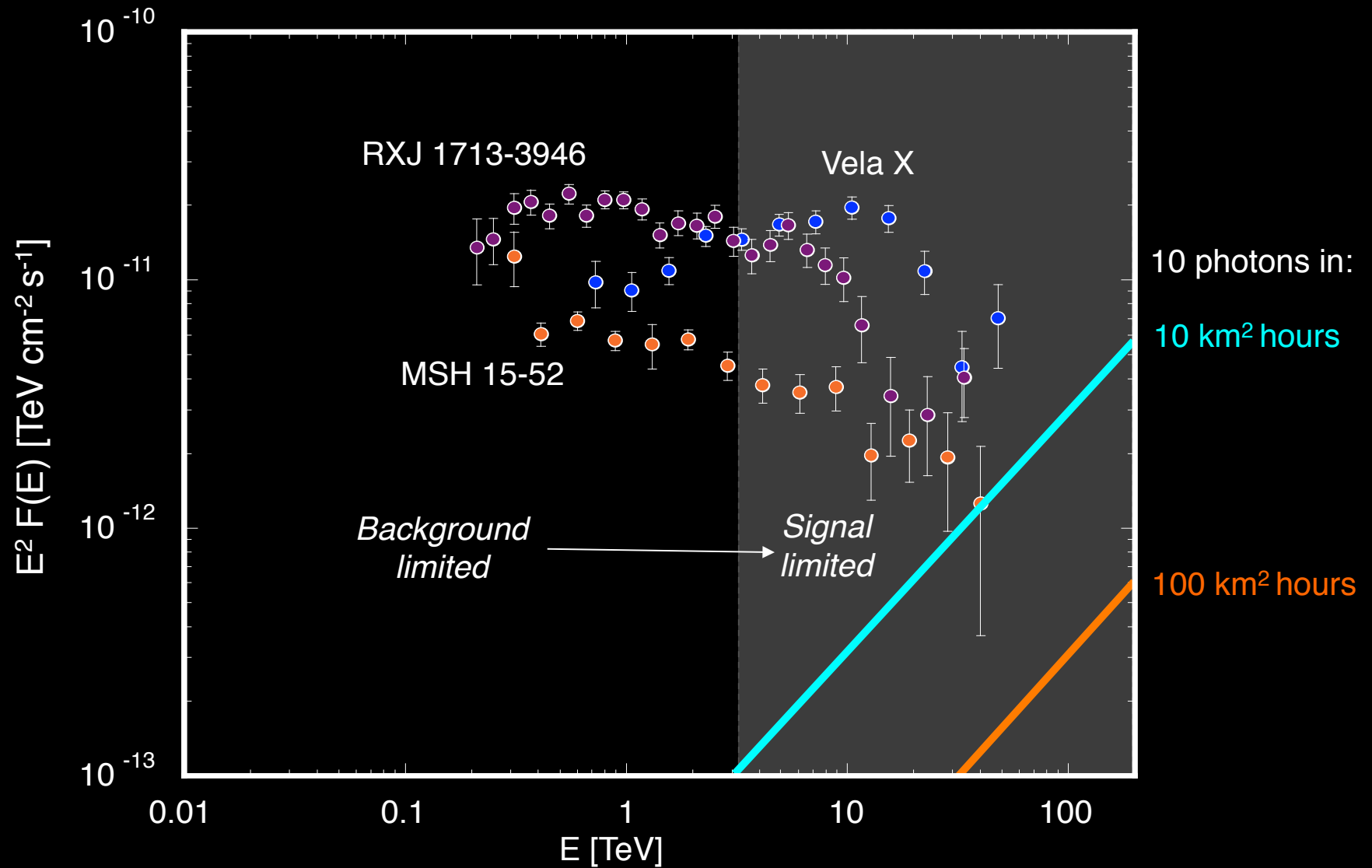
THE CHERENKOV TELESCOPE ARRAY



CTA GOAL SENSITIVITY

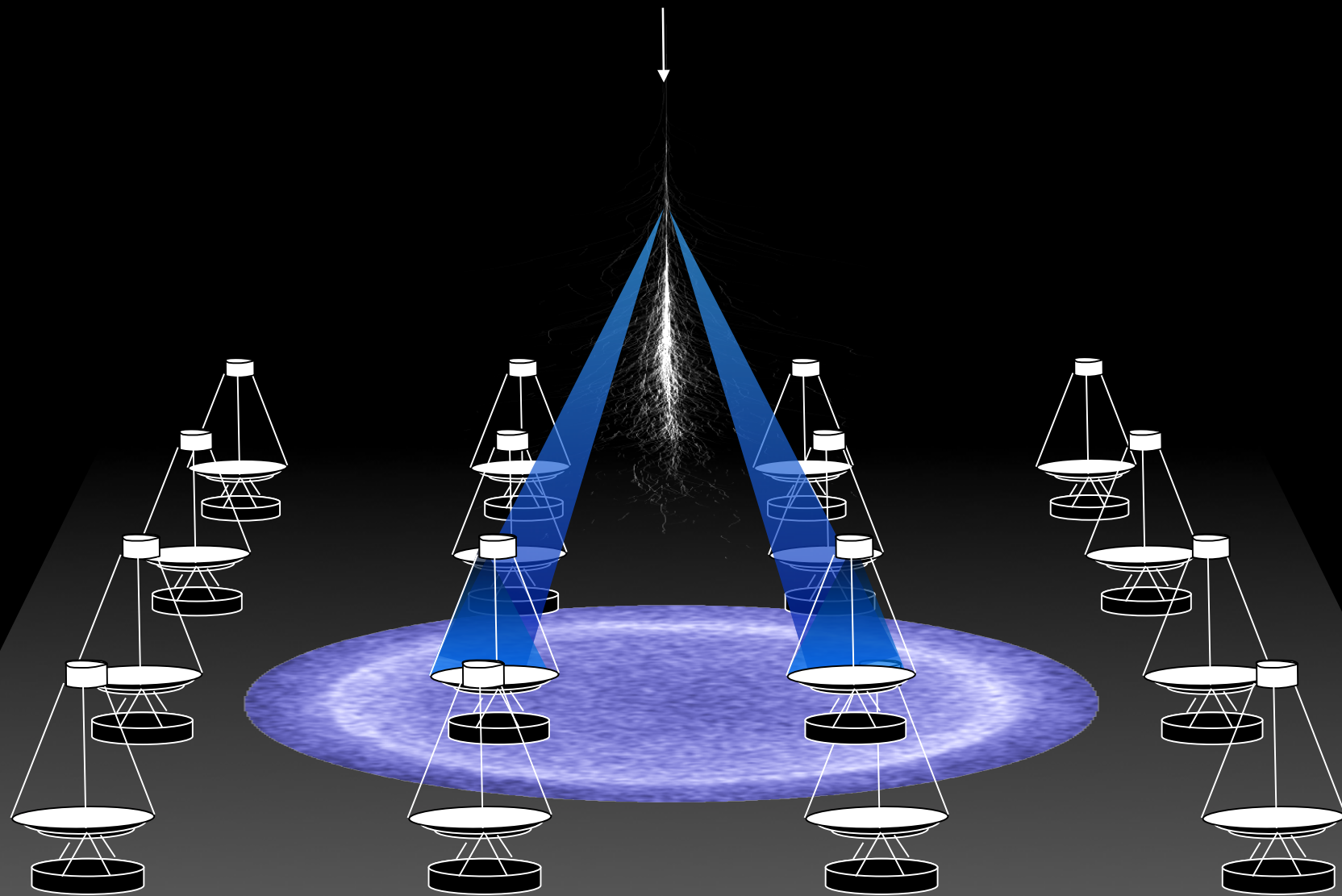


THE PROBLEM AT HIGH ENERGIES



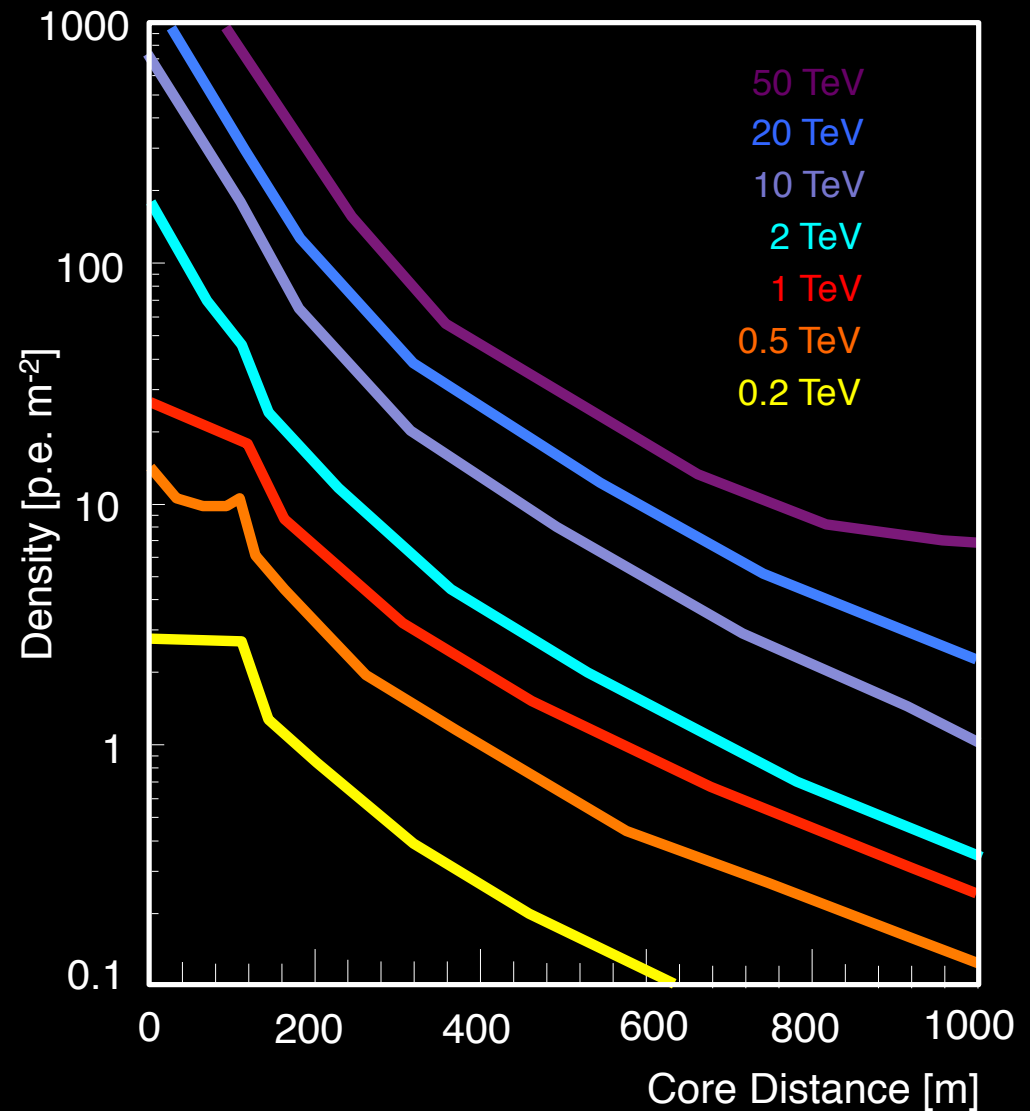
Adapted from Werner Hofmann

COLLECTING CHERENKOV LIGHT



TELESCOPE SPACING

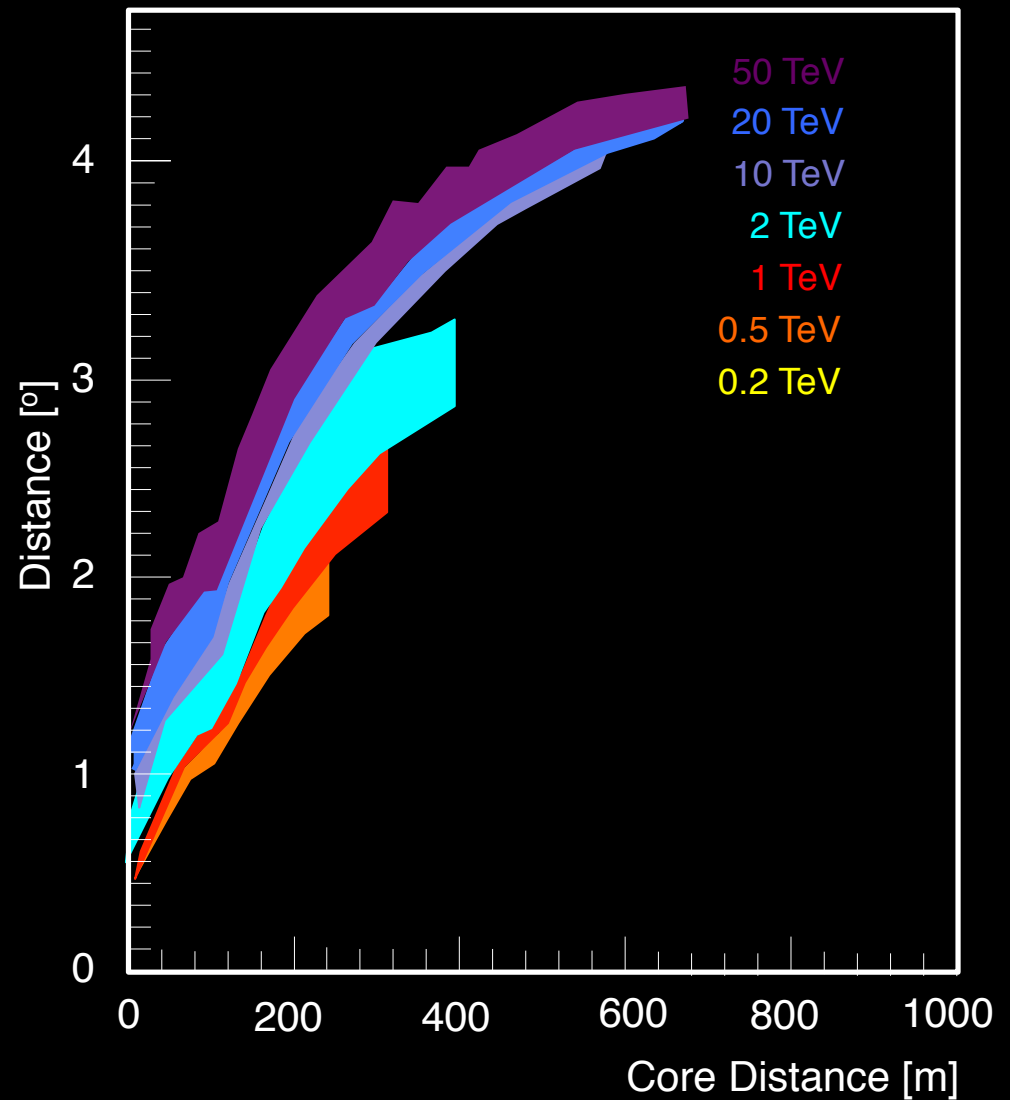
- ⊙ Cherenkov light useful out to >600 m at few TeV
- ⊙ Greater spacing requires larger reflectors.



Adapted from de La Calle Perez, Biller, 2006

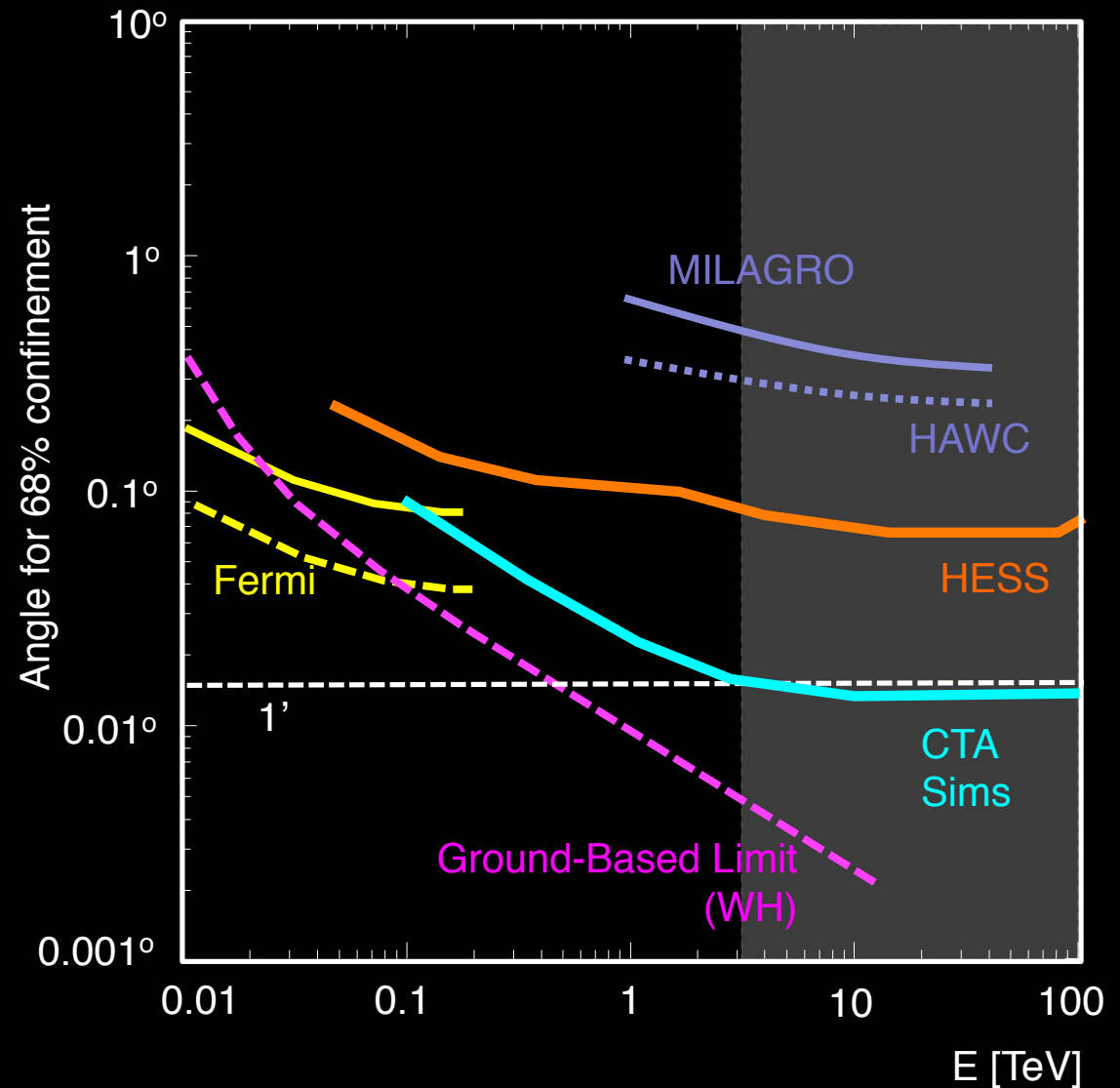
FIELD OF VIEW

- ⊙ Cherenkov light useful out to >600 m at few TeV
- ⊙ Greater spacing requires larger reflectors.
- ⊙ Greater spacing requires larger field of view:
 - 3 - 5° at 100 m (current)
 - 7 - 8° at 200 m
 - 10 - 11° at 600 m



THE OPPORTUNITY AT HIGH ENERGIES

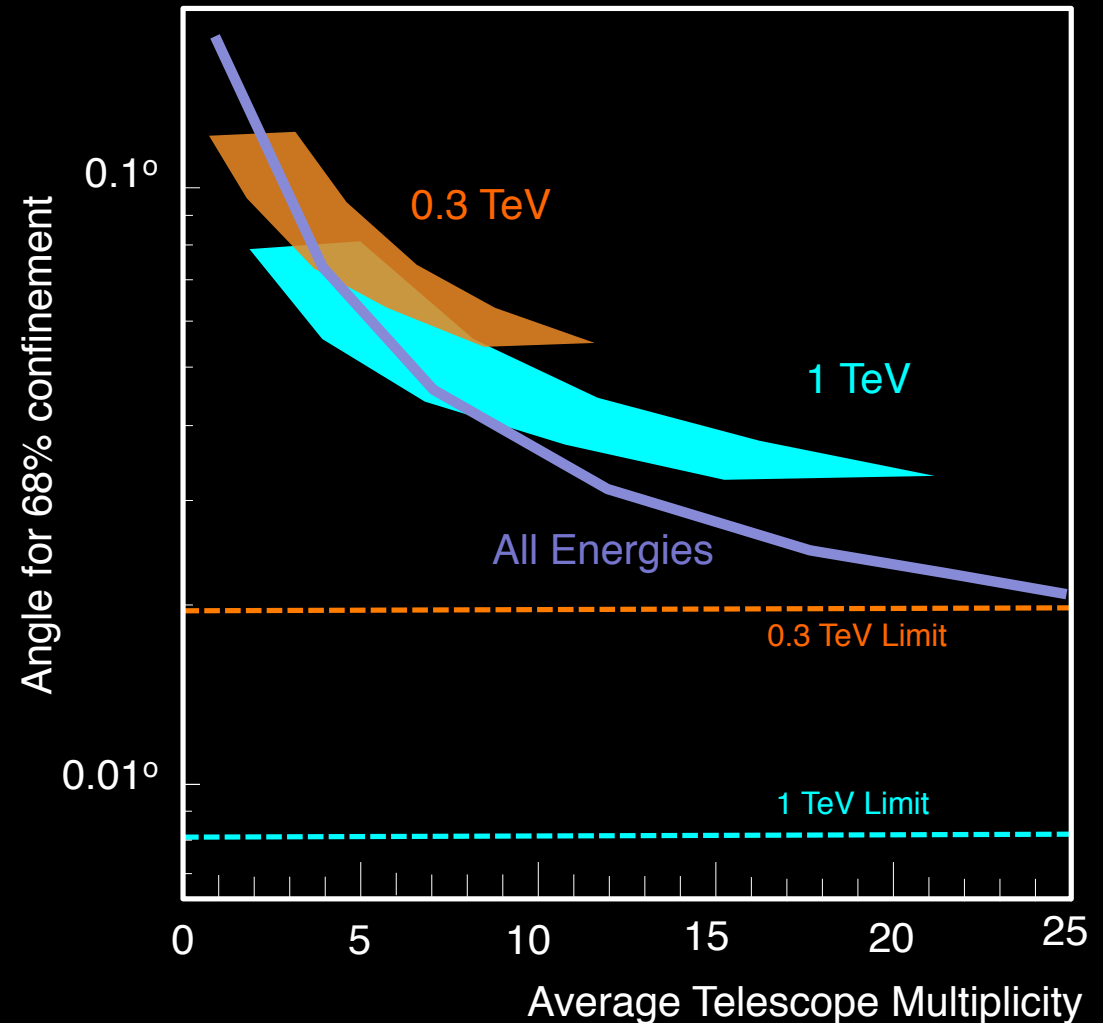
- ⊙ Angular Resolution!
- ⊙ High precision only achievable $> \text{TeV}$:
 - Limit: few " at 30 TeV
 - 1' is achievable at 100 TeV with modest collection area.
 - x 5 improvement from HESS



Adapted from Funk, Reimer, Torres, Hinton 2008

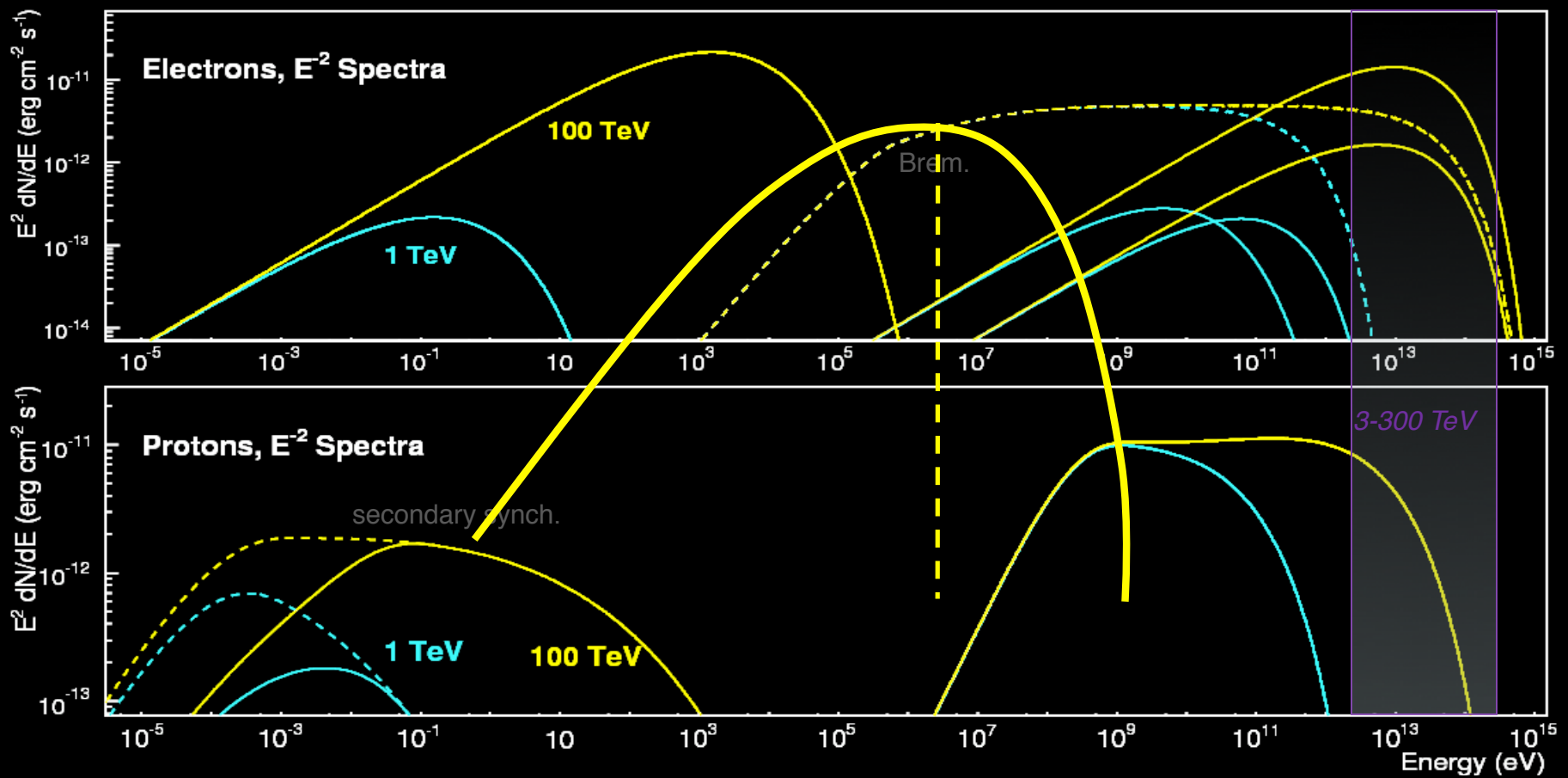
THE OPPORTUNITY AT HIGH ENERGIES

- ⊙ Angular Resolution!
- ⊙ High precision only achievable $> \text{TeV}$:
 - Limit: few " at 30 TeV
 - 1' is achievable at 100 TeV with modest collection area.
 - x 5 improvement from HESS
- ⊙ Angular Res. improves with multiplicity:
 - At 1 TeV ~16 telescopes.
 - $> 1 \text{ TeV}$ a very high multiplicity is needed for high angular resolution



Adapted from Funk & Hinton 2009

WHAT ARE WE LOOKING FOR?



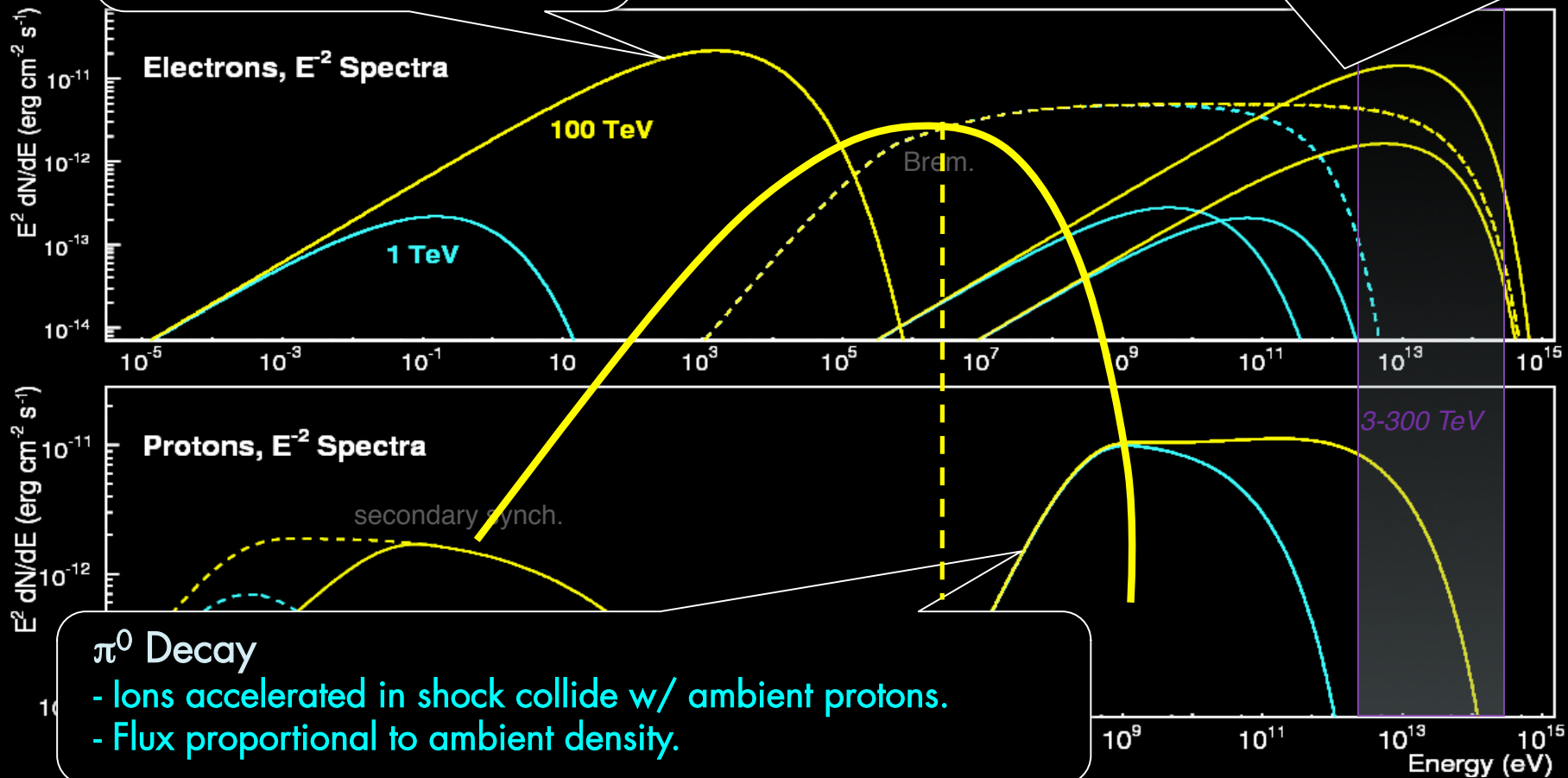
WHAT ARE WE LOOKING FOR?

X-Ray Synchrotron (keV)

- Accelerated e- in B Field
- Depends on B Field

Gamma-Ray Inverse Compton (TeV)

- Acc. e- upscatter ambient photons to γ -rays
- CMB, local dust, starlight provide seed photons.

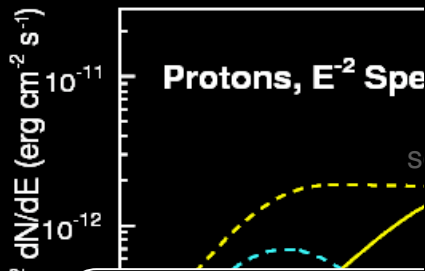
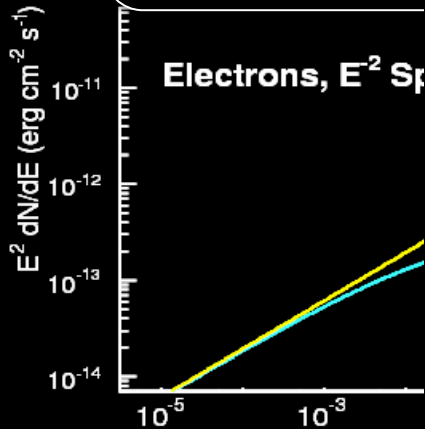


WHAT ARE WE LOOKING FOR?

X-Ray Synchrotron (keV)

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$$x^2 F_\gamma(x, E_p)$$

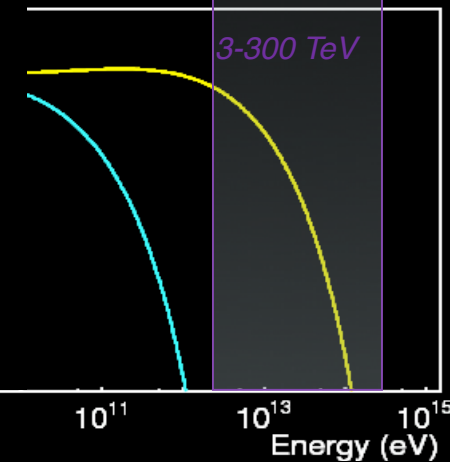
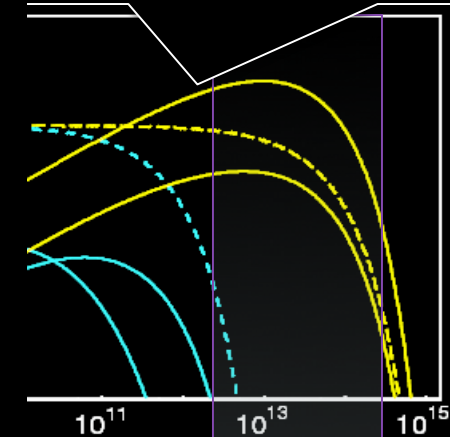
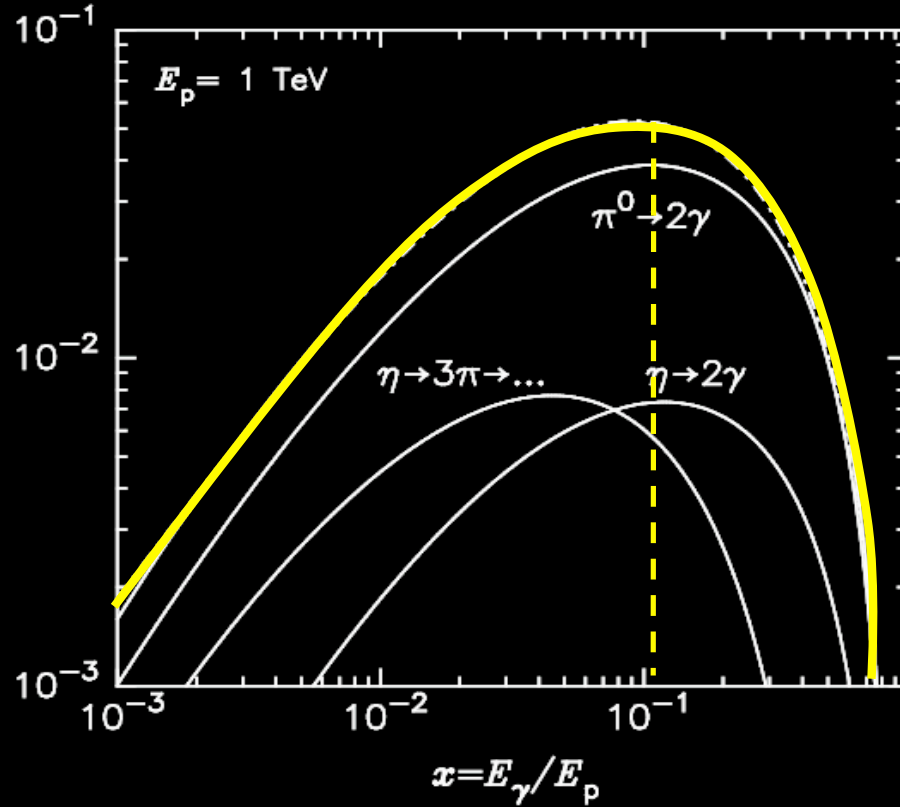


π^0 Decay

- Ions accelerated
- Flux proportional to ambient density.

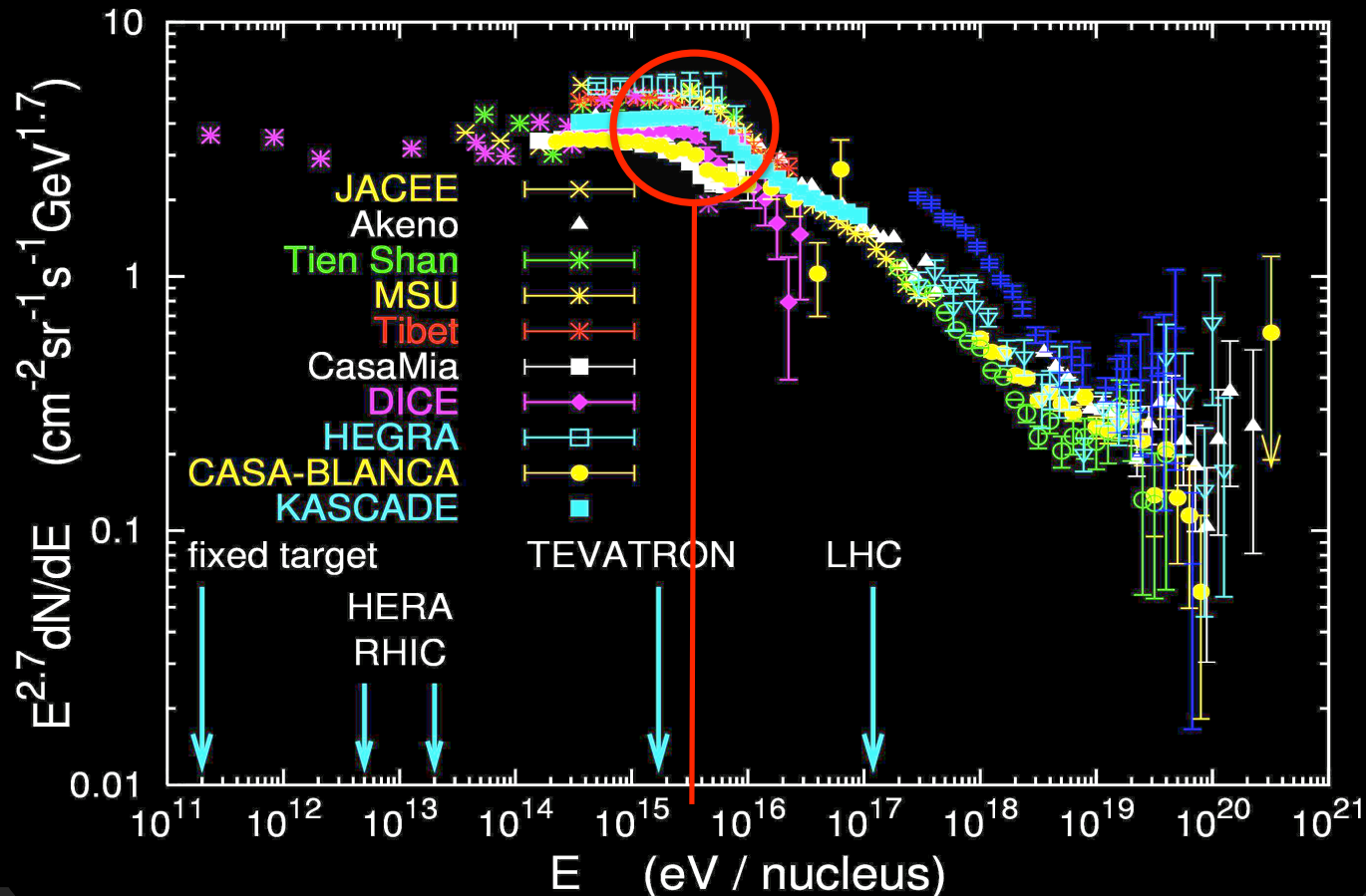
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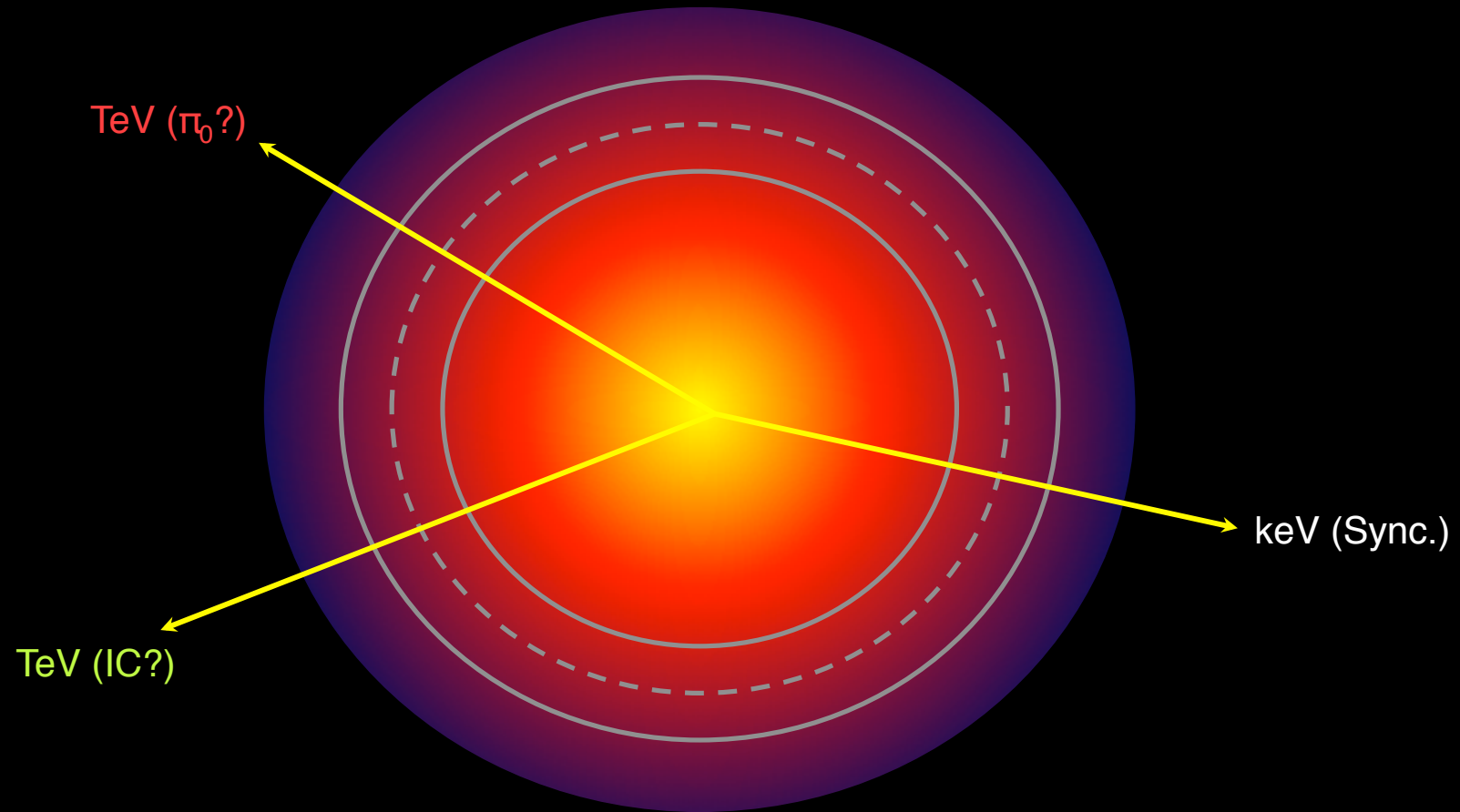


GALACTIC PEVATRONS

- ◉ Cosmic-Ray spectrum is smooth until ~ 3 PeV
 - Galactic origin at least this far
 - SNR are a good candidate...

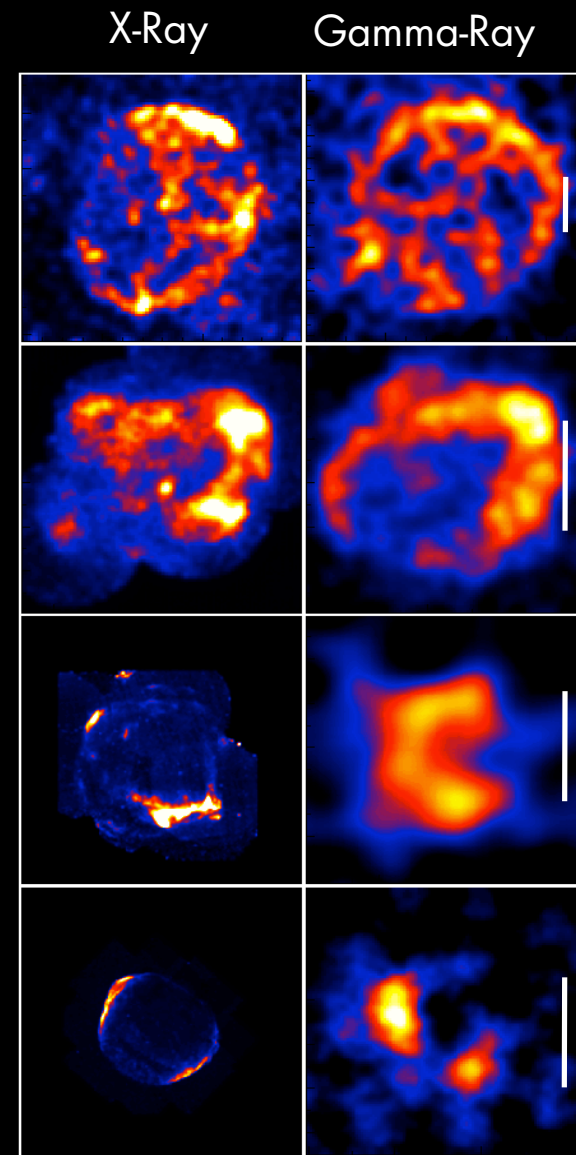


MORPHOLOGY: 1990s



MORPHOLOGY: 2000s

- ⊙ 4 shell-like objects detected in TeV.
- ⊙ Young historical SNRs:
 - RX J1713.7-3946
 - Vela Junior
 - RCW 86
 - SN 1006
- ⊙ Point-source emission from Cas A.
- ⊙ All resolved shells show correlation with X-Ray emission.
- ⊙ Electron acceleration to > 100 TeV



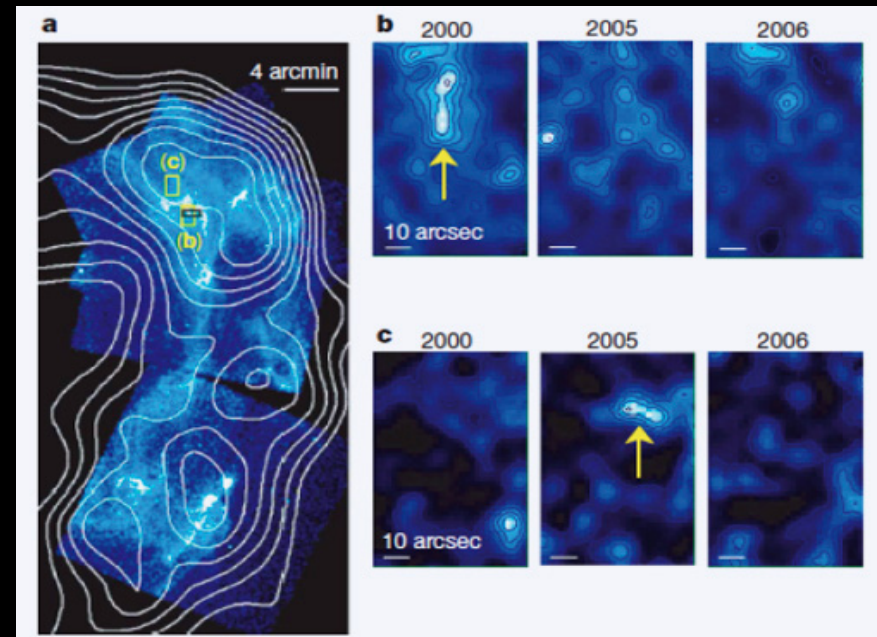
MORPHOLOGY: 2000s

◎ B-Field Amplification in SNR

- Evidence:
 - X-Ray Filaments
 - X-Ray Variability
- Implies CR Pressure is significant
- $<10\%$ radius B-Field structure expected.

◎ TeV Morphology provides:

- Differentiate between leptonic and hadronic acceleration
- Test theory of magnetic field amplification in CR modified shocks.



Reproduced from Uchiyama et al., 2007

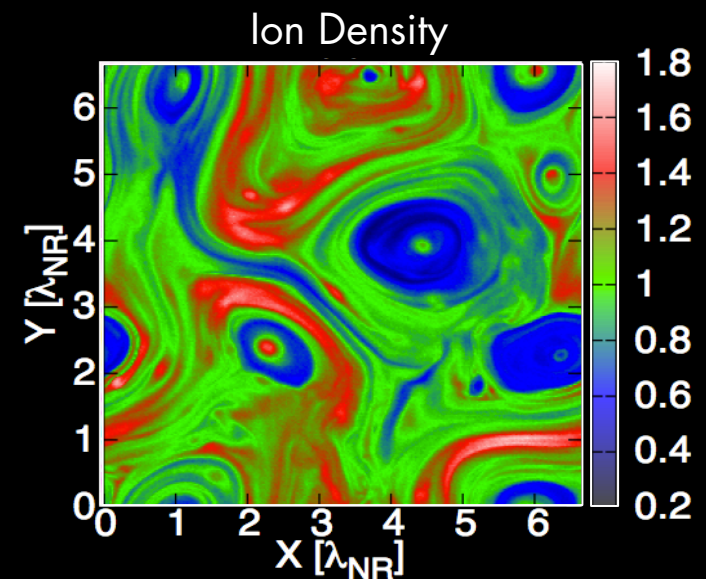
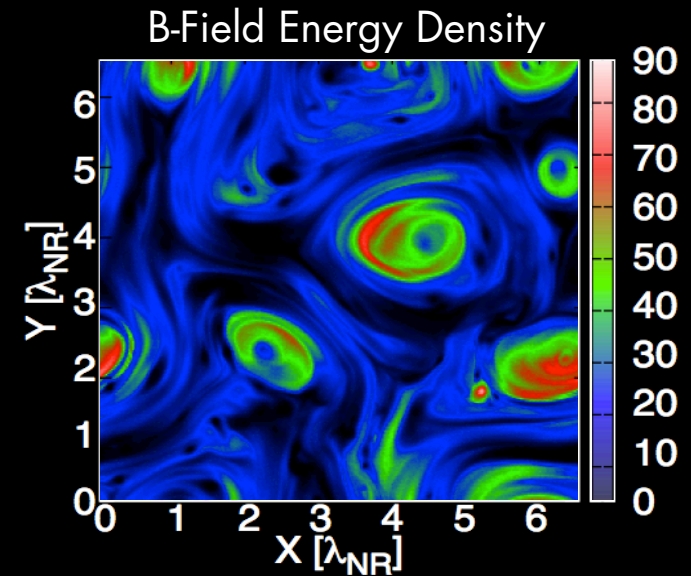
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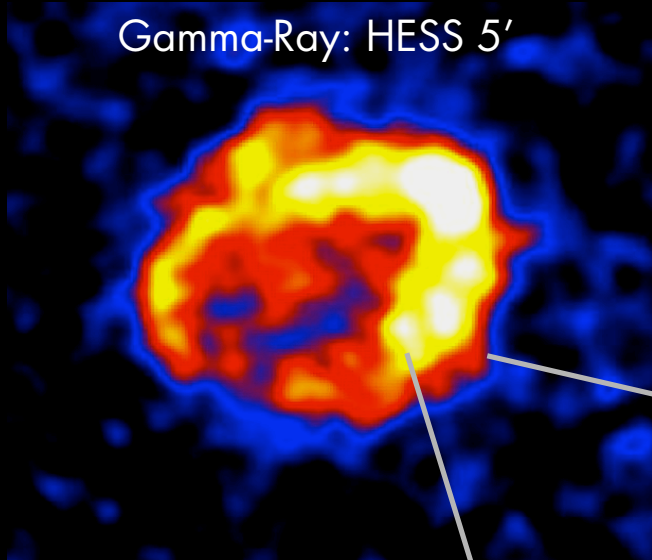


Reproduced from Ohira, Reville, Kirk 2009

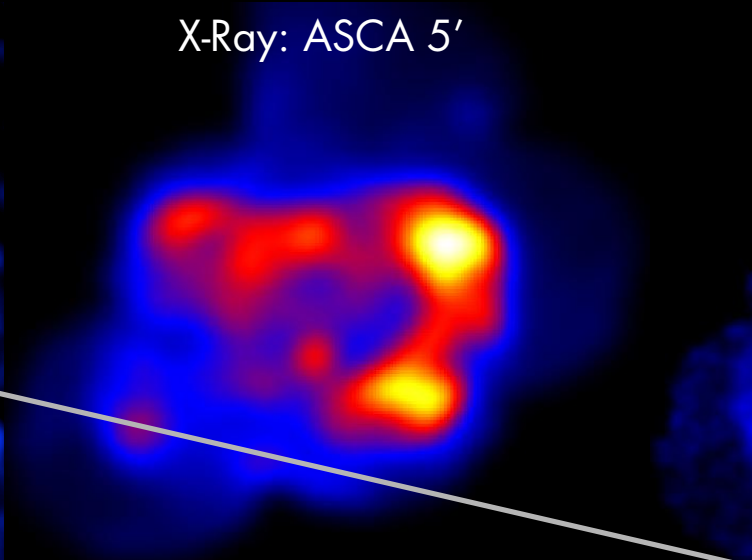
MORPHOLOGY: 2000s

RX J1713.7-3946

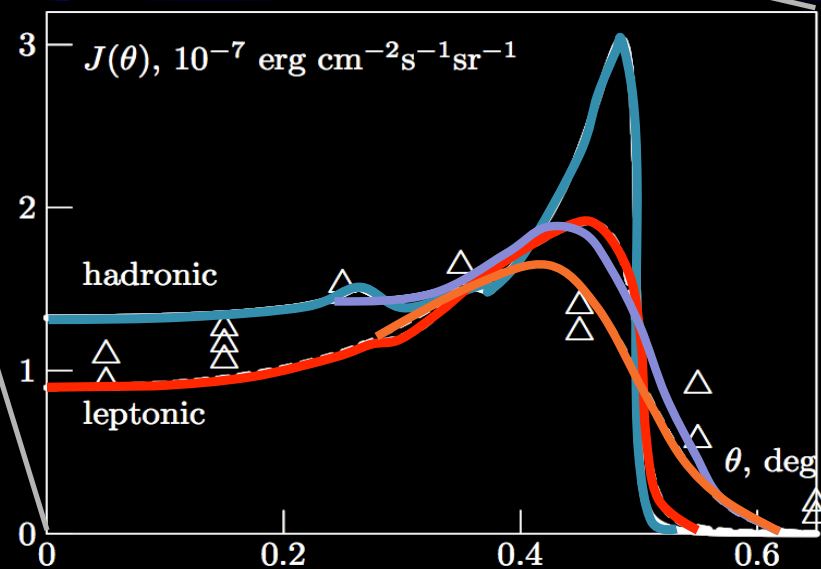
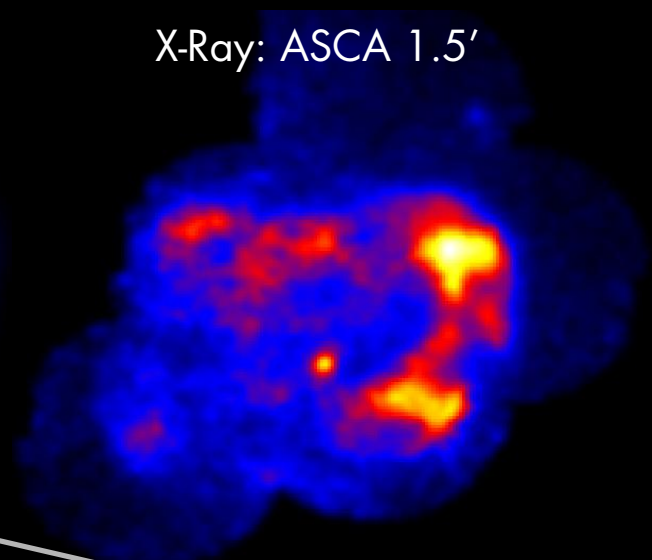
Gamma-Ray: HESS 5'



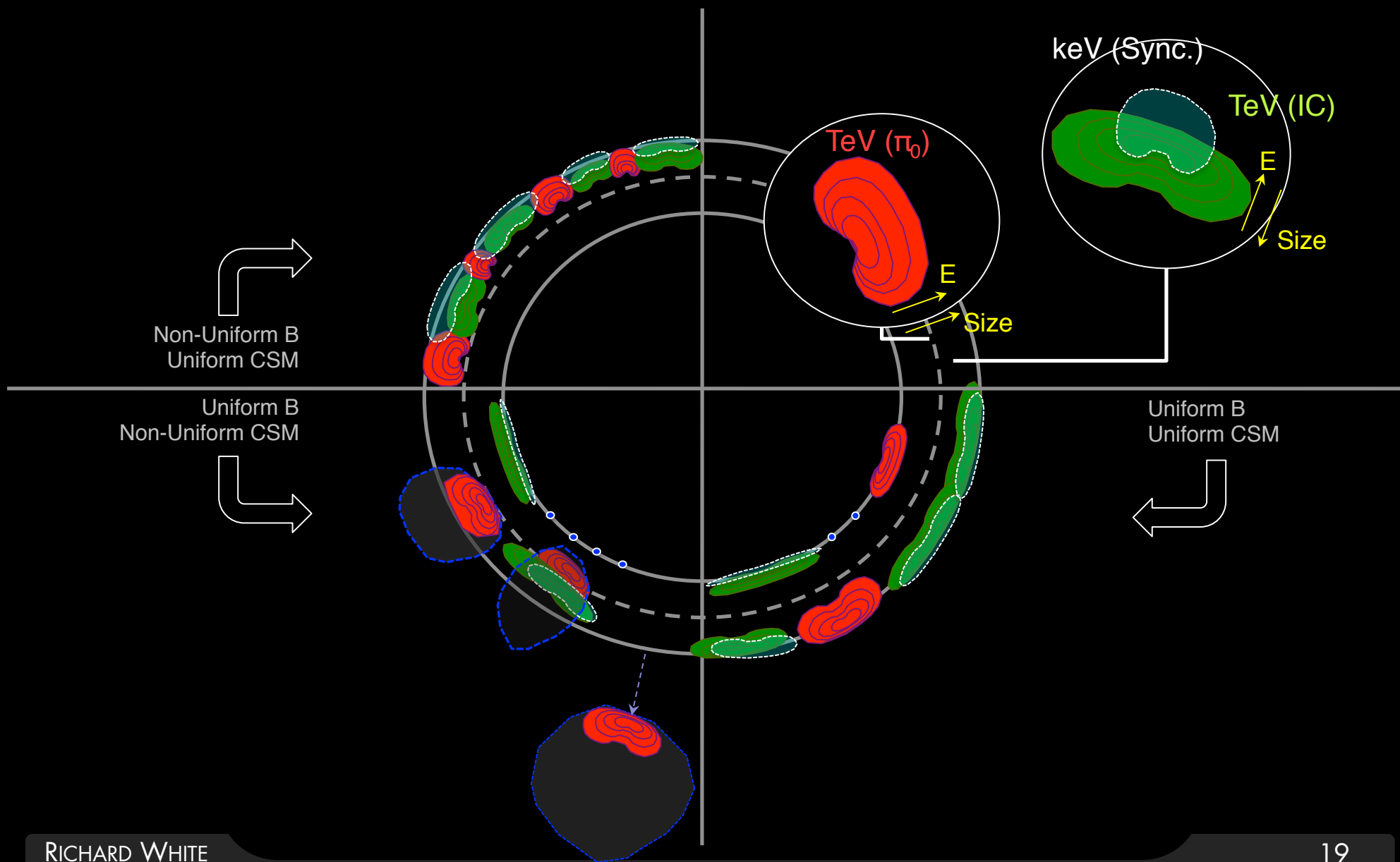
X-Ray: ASCA 5'



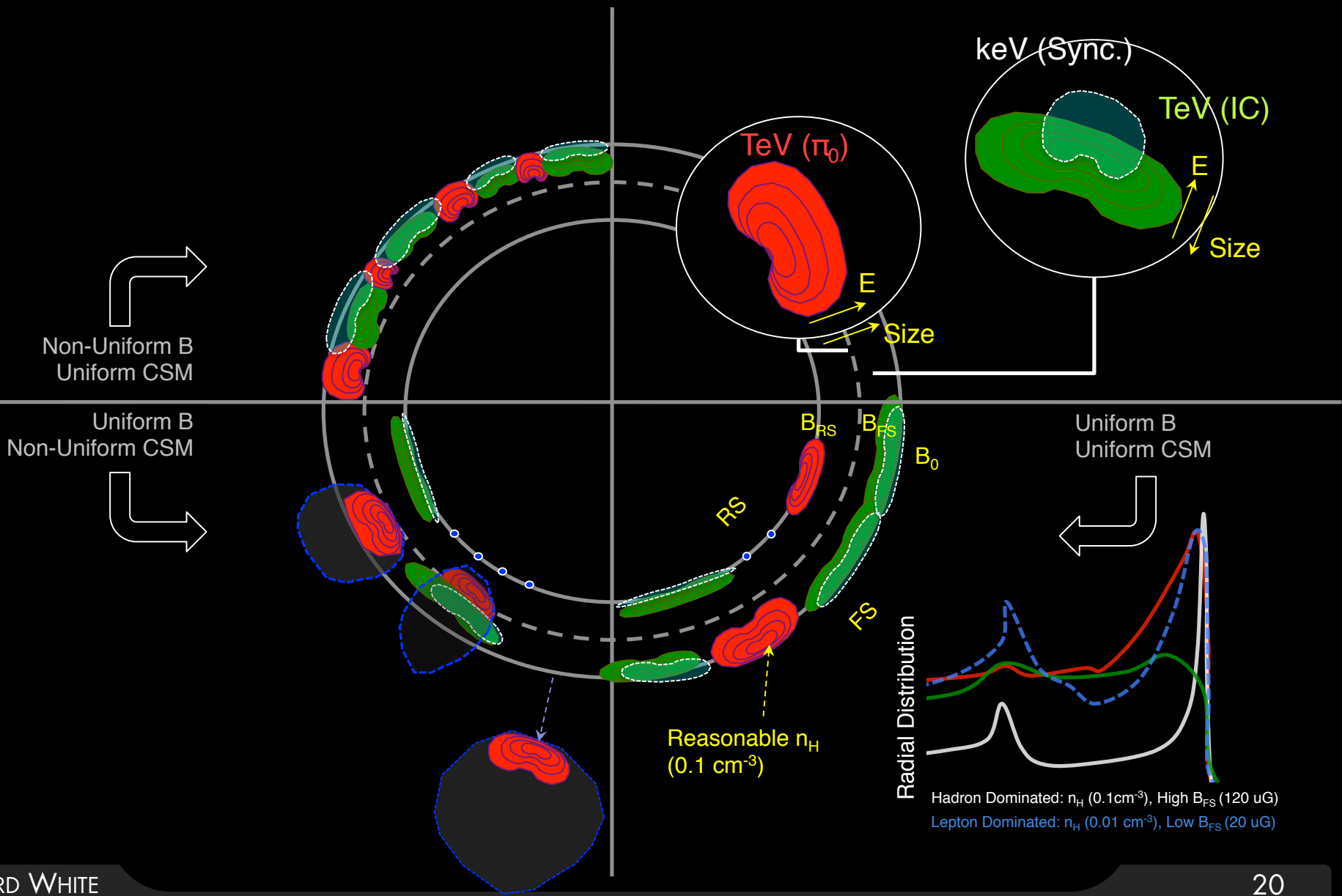
X-Ray: ASCA 1.5'



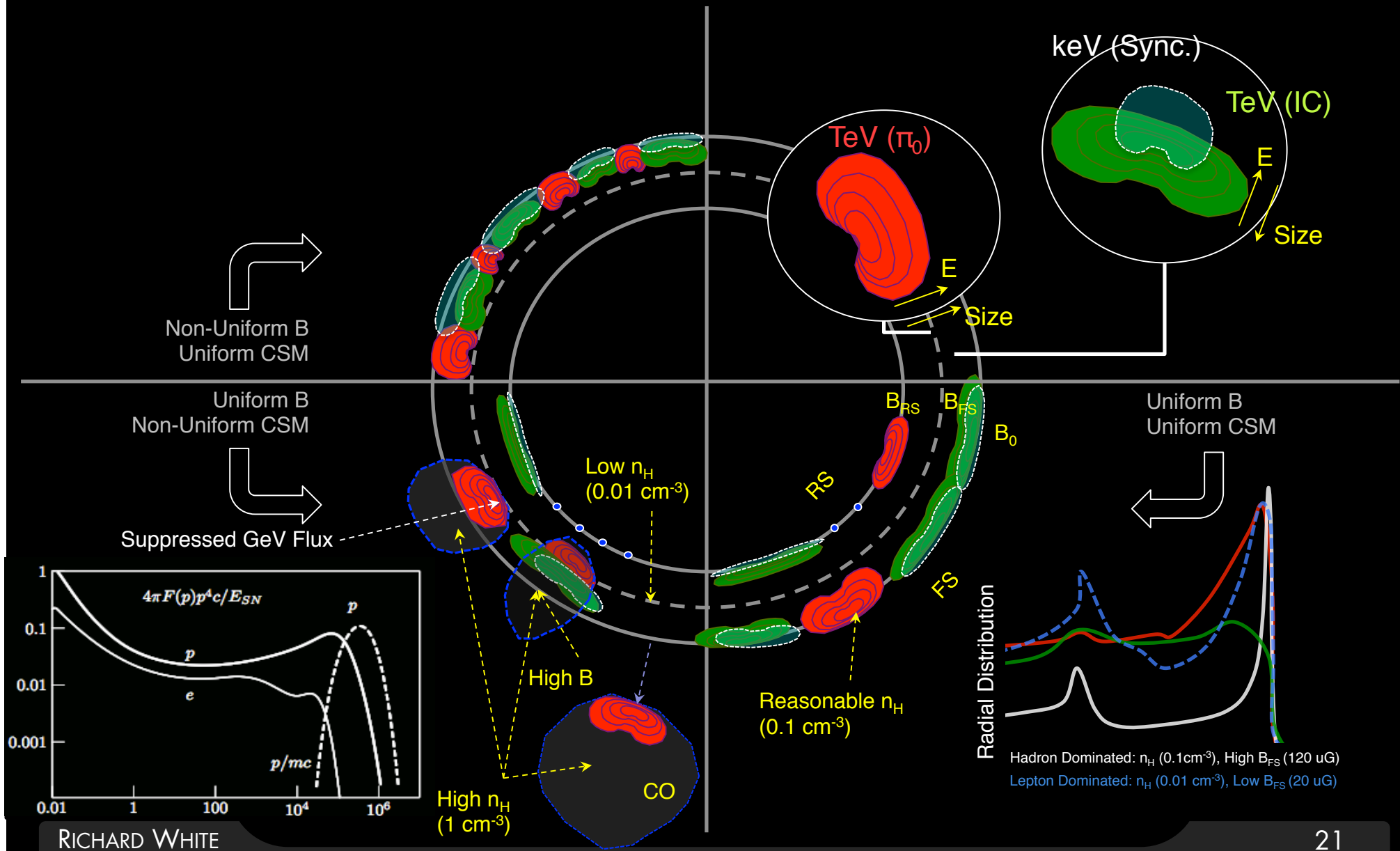
MORPHOLOGY: 2010s



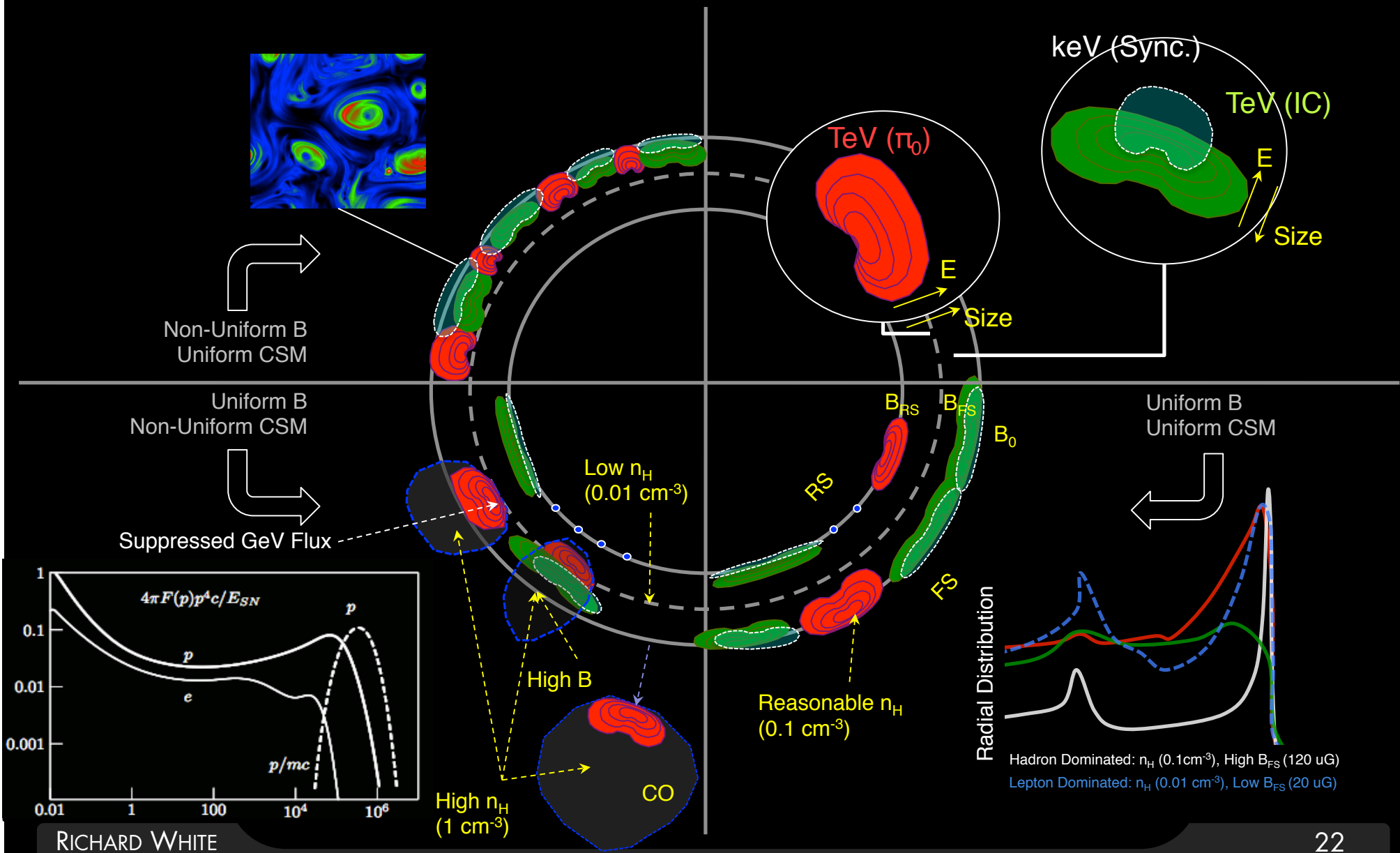
MORPHOLOGY: 2010s



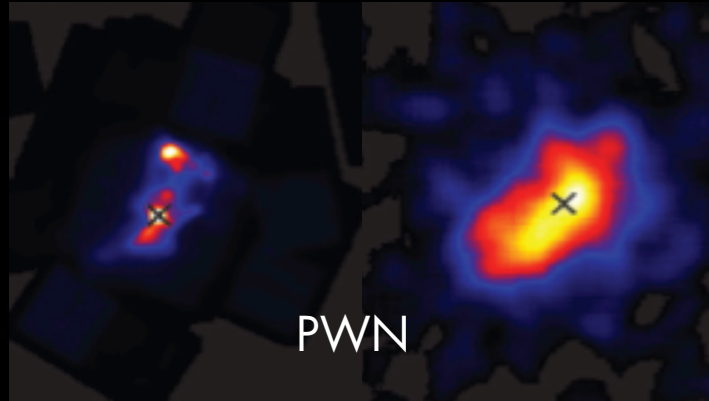
MORPHOLOGY: 2010s



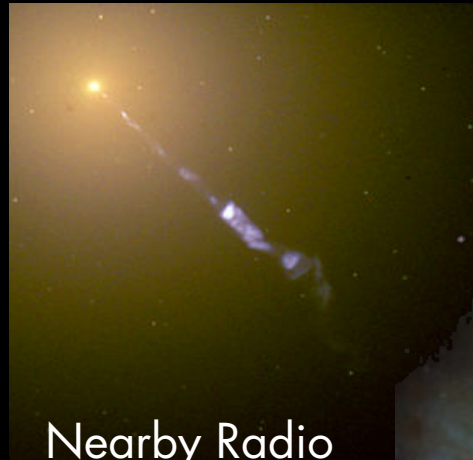
MORPHOLOGY: 2010s



WHAT ELSE:



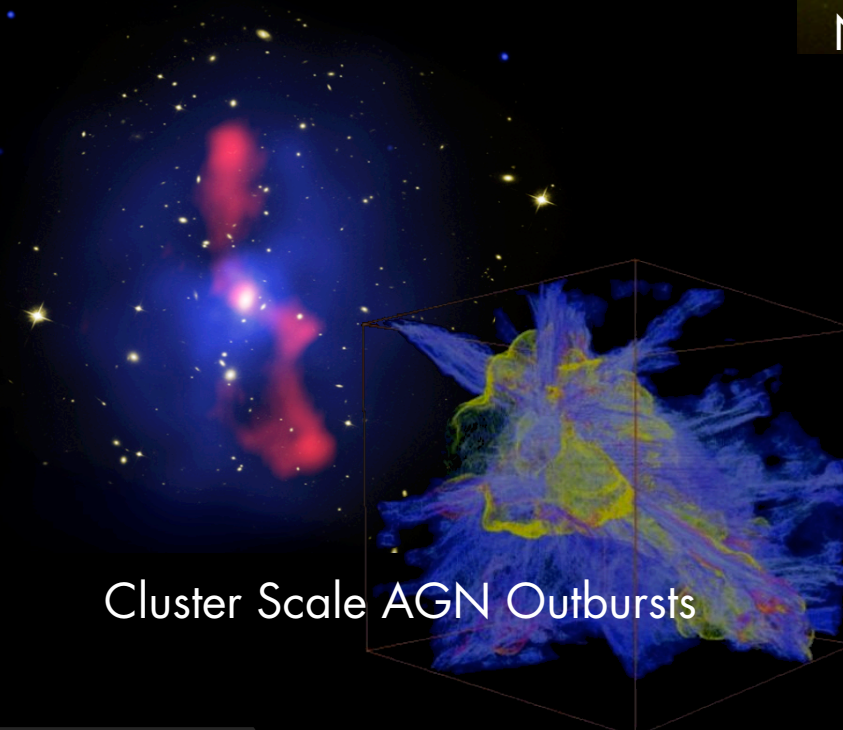
PWN



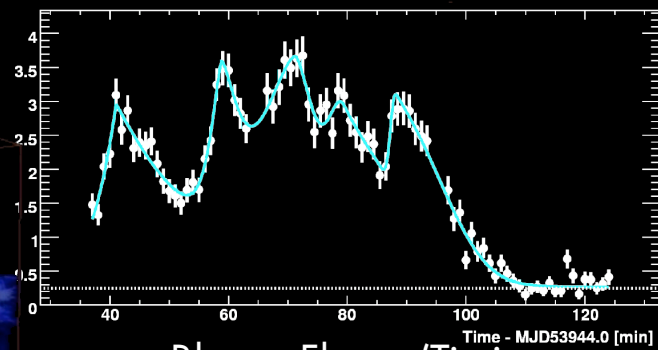
Nearby Radio Galaxies



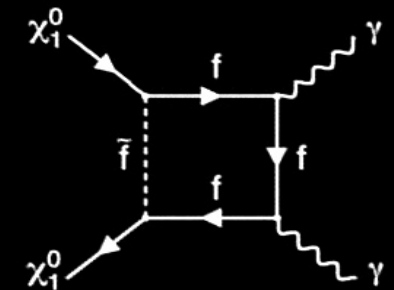
Starburst Galaxies



Cluster Scale AGN Outbursts



Blazar Flares (Timing, Lorentz Invariance)



Decay of Heavy WIMPs

CONCLUSIONS

- ⊙ CTA will provide unprecedented sensitivity across a wide energy range.
- ⊙ 3 – 300 TeV sensitivity is critical to address major questions in astroparticle physics.
- ⊙ Excellent angular resolution (the best possible anywhere above ~ 100 keV) is possible and **required**.